

General Relativistic Radiation MHD Simulations of Black Hole Accretion

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Students:
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The Kavli Institute for
Theoretical Physics
University of California, Santa Barbara

- Long-standing questions in black hole astrophysics
 - Are radiation-pressure dominated accretion disks thermally unstable? (Pringle, Rees & Pacholczyk 1973)
 - Are radiation-pressure dominated accretion disks secularly unstable? (Lightman & Eardley 1974)
 - Can a disk sustain locally super-Eddington fluxes? (Begelman 2001)
 - How are the jets associated with black holes affected when the disk is radiation-pressure dominated?
- Building supermassive black holes
- Simulating quasi-stars



M87

- GRMHD simulations + ray-tracing codes

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 - Schnittman et al. (2006)
 - Noble et al. (2007); Mościbrodzka et al. (2009)
 - Dexter et al. (2009, 2010)

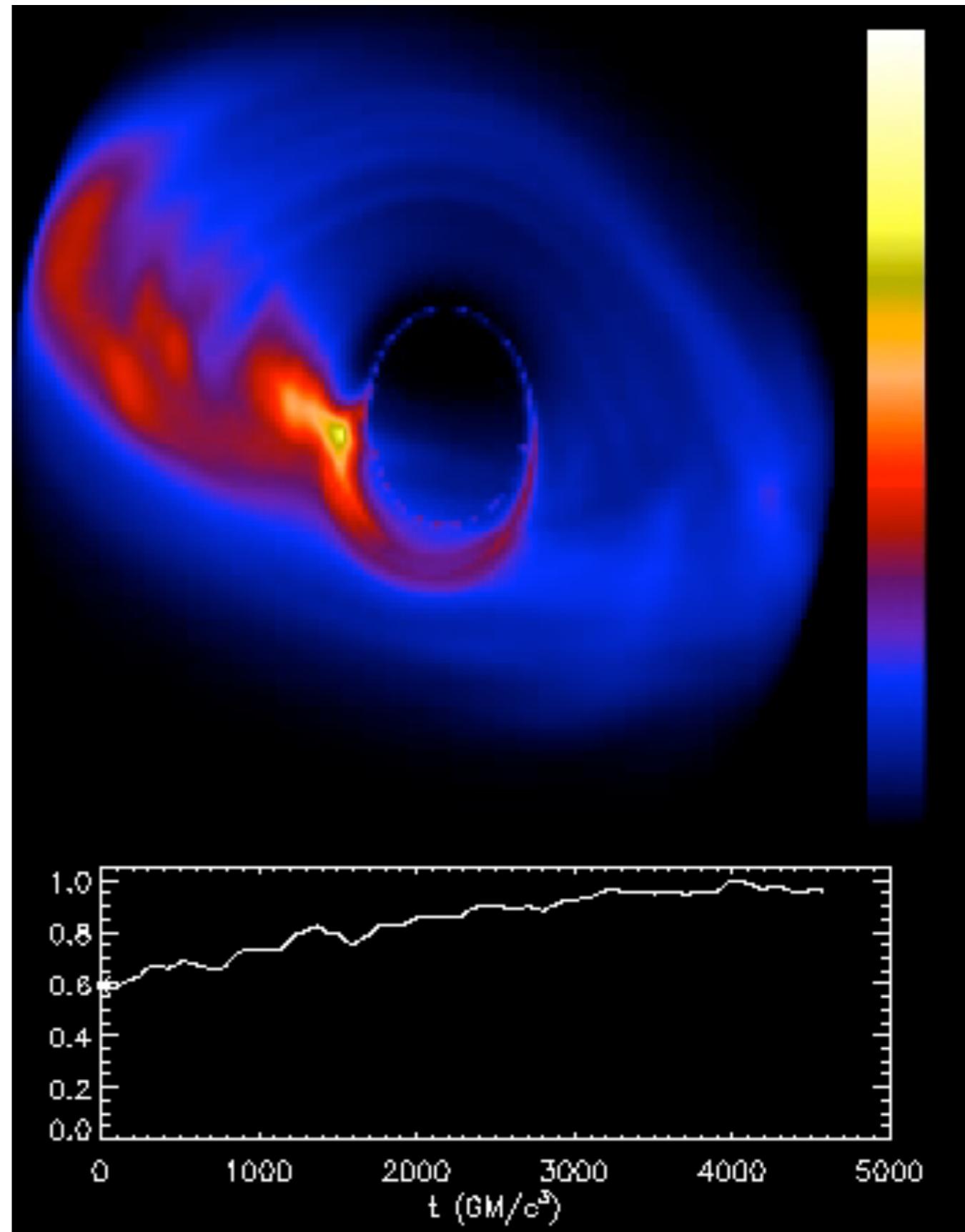
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- Applicable to very low-luminosity systems
 - Sgr A*
 - M87 (maybe)

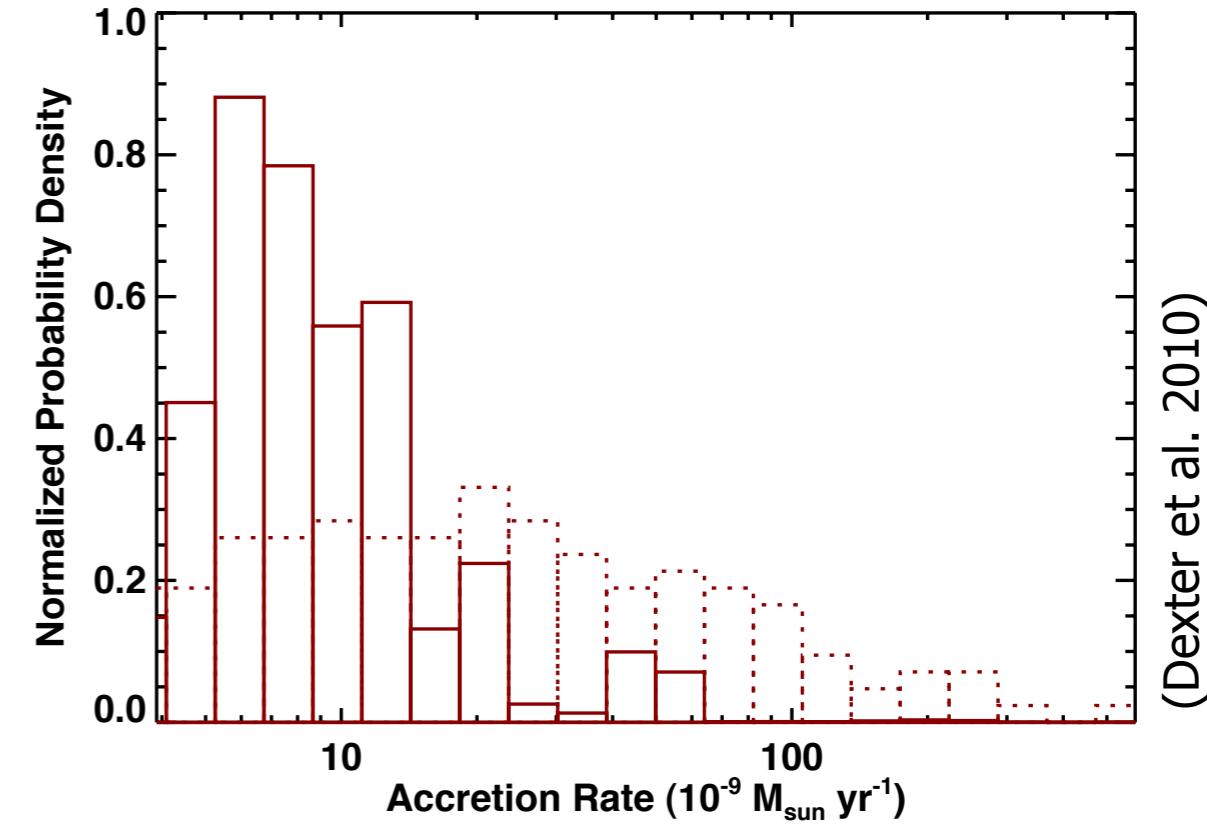
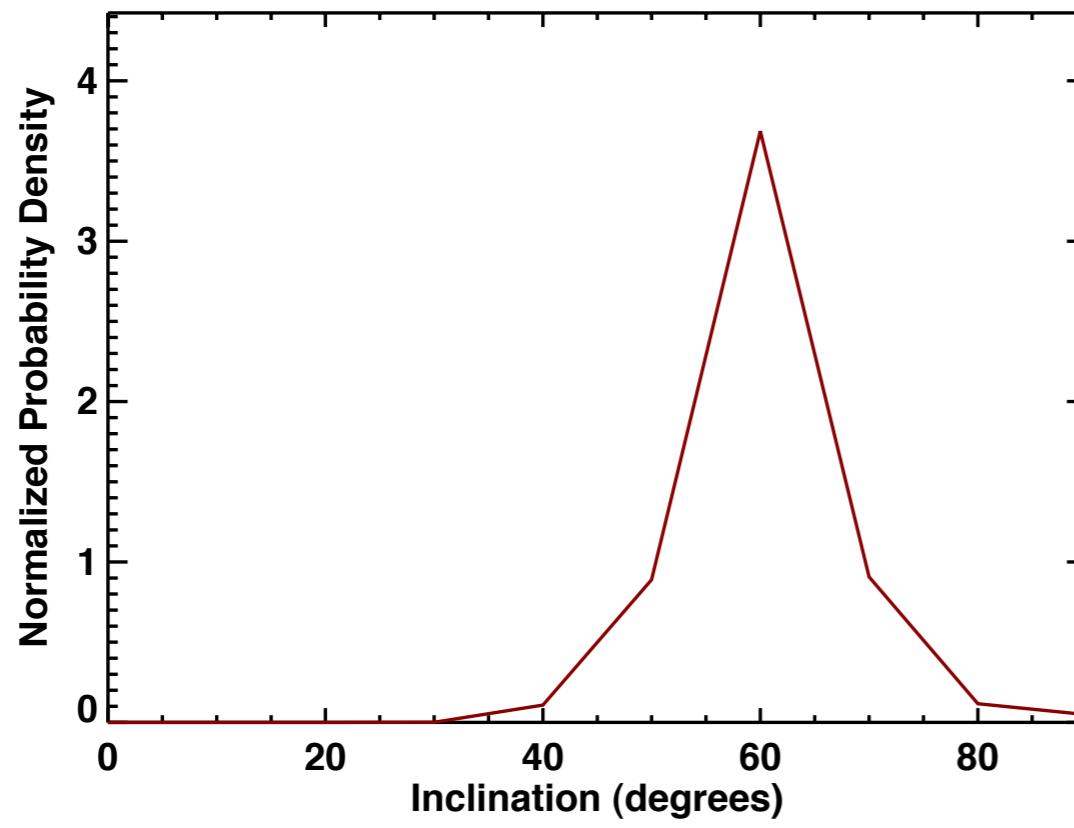
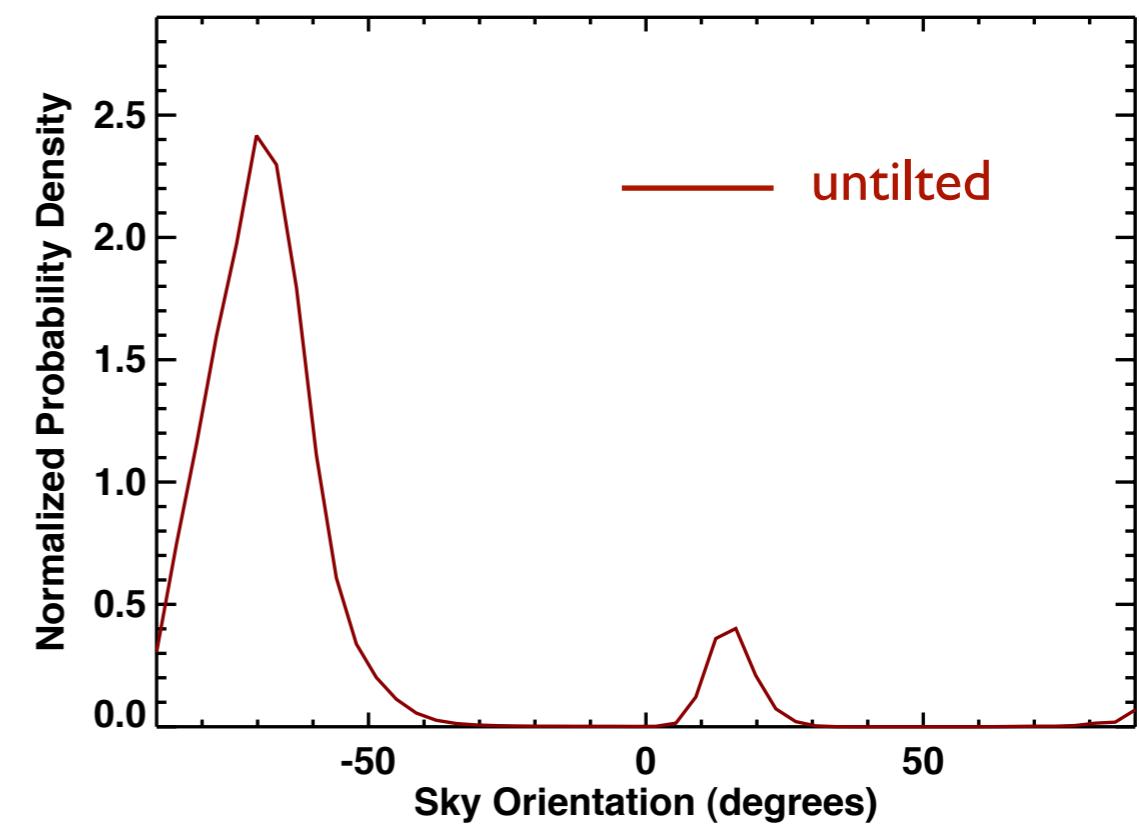
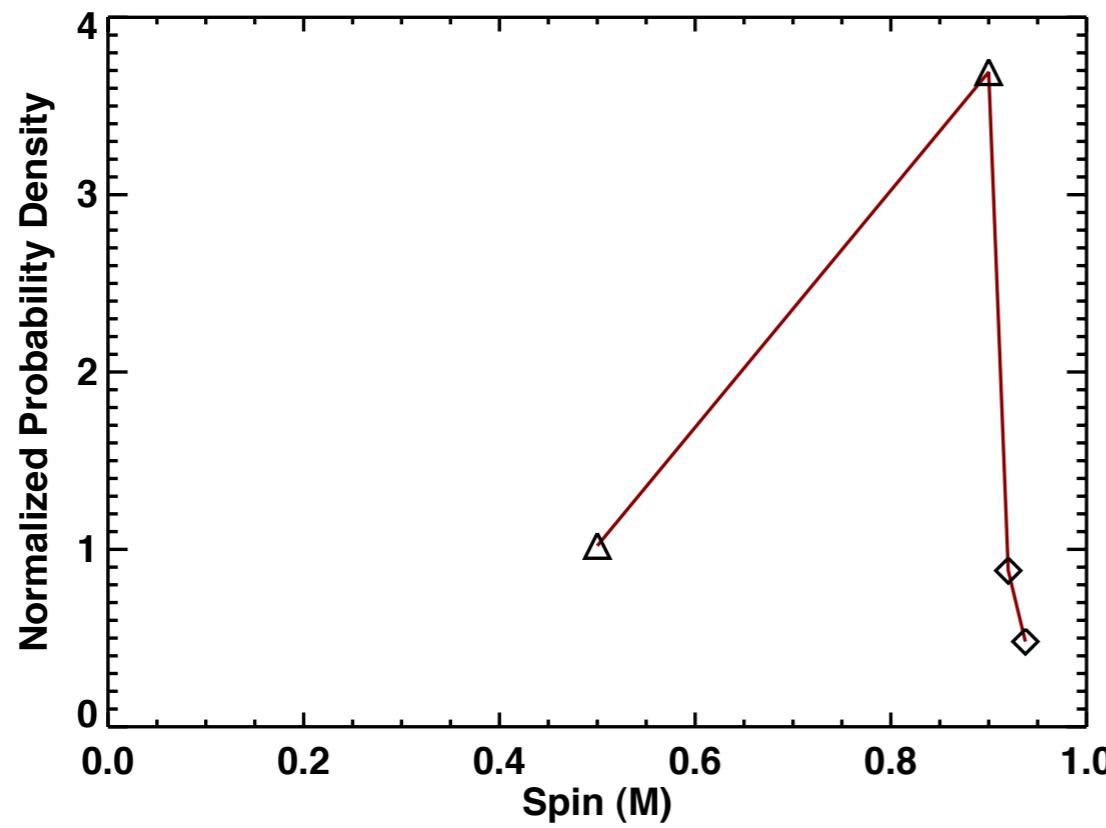
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Credit: Jason Dexter

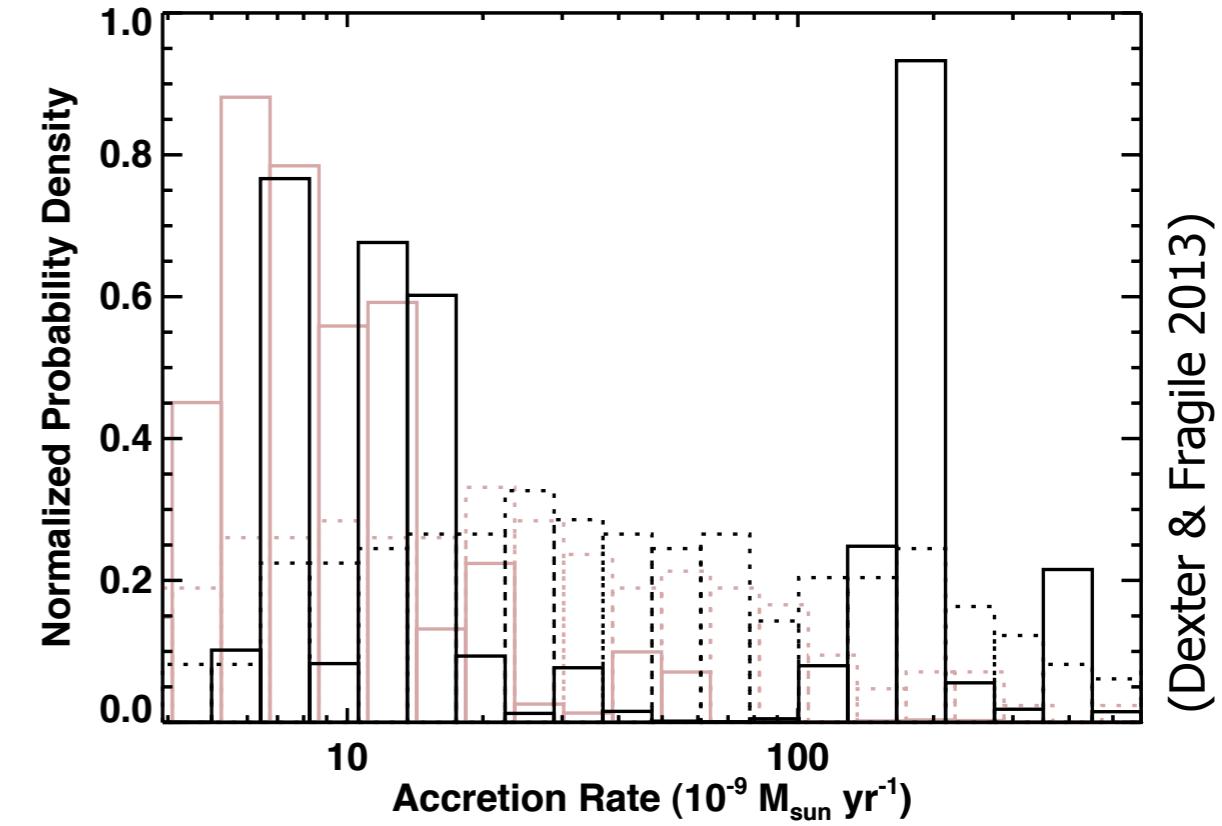
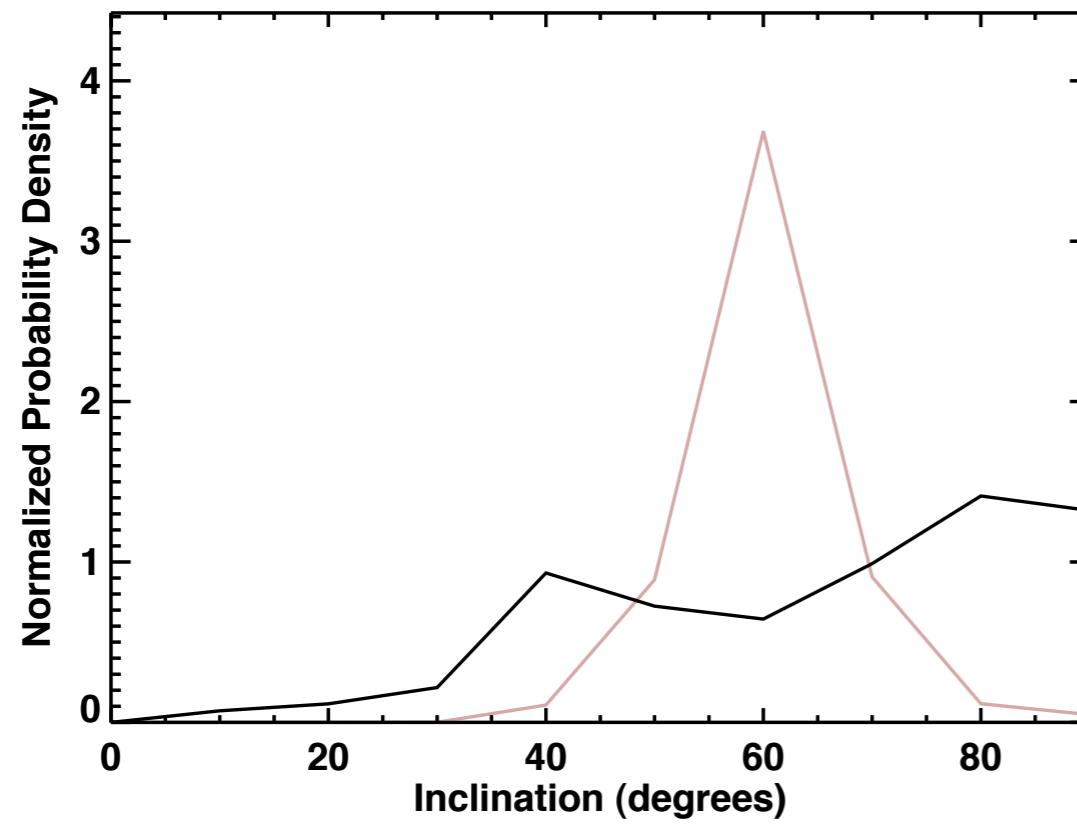
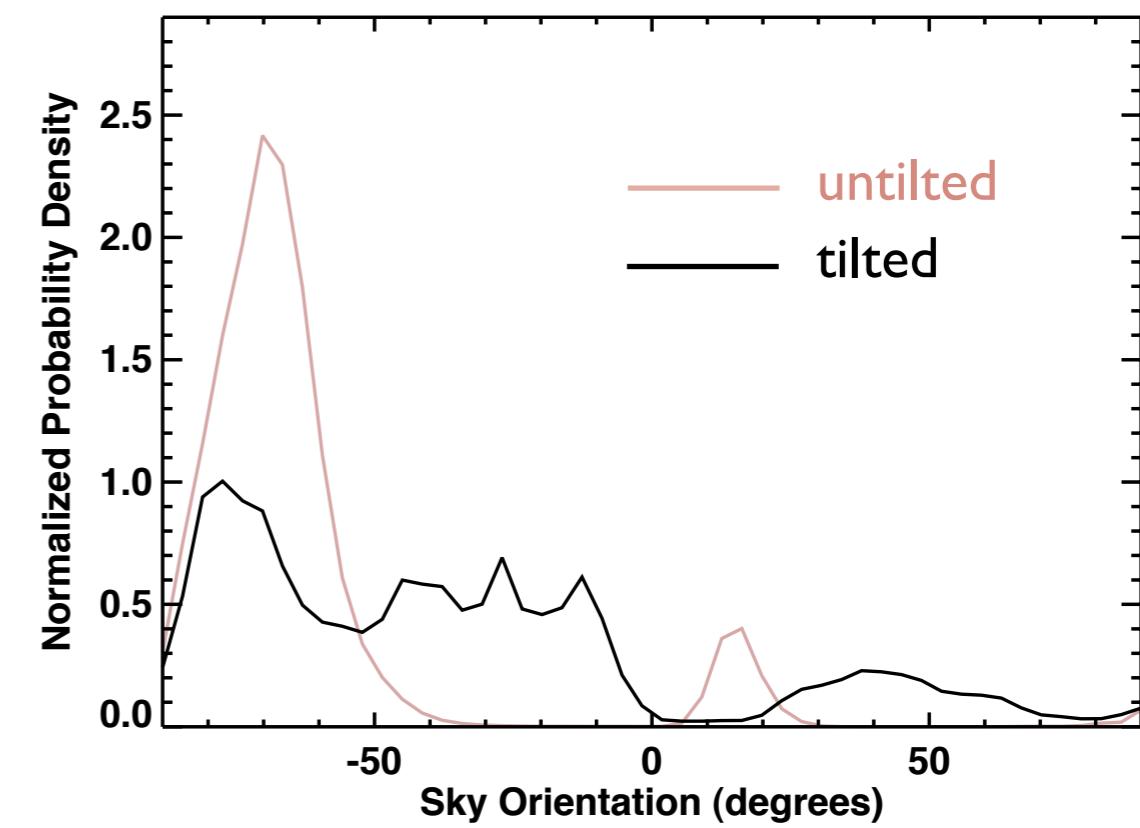
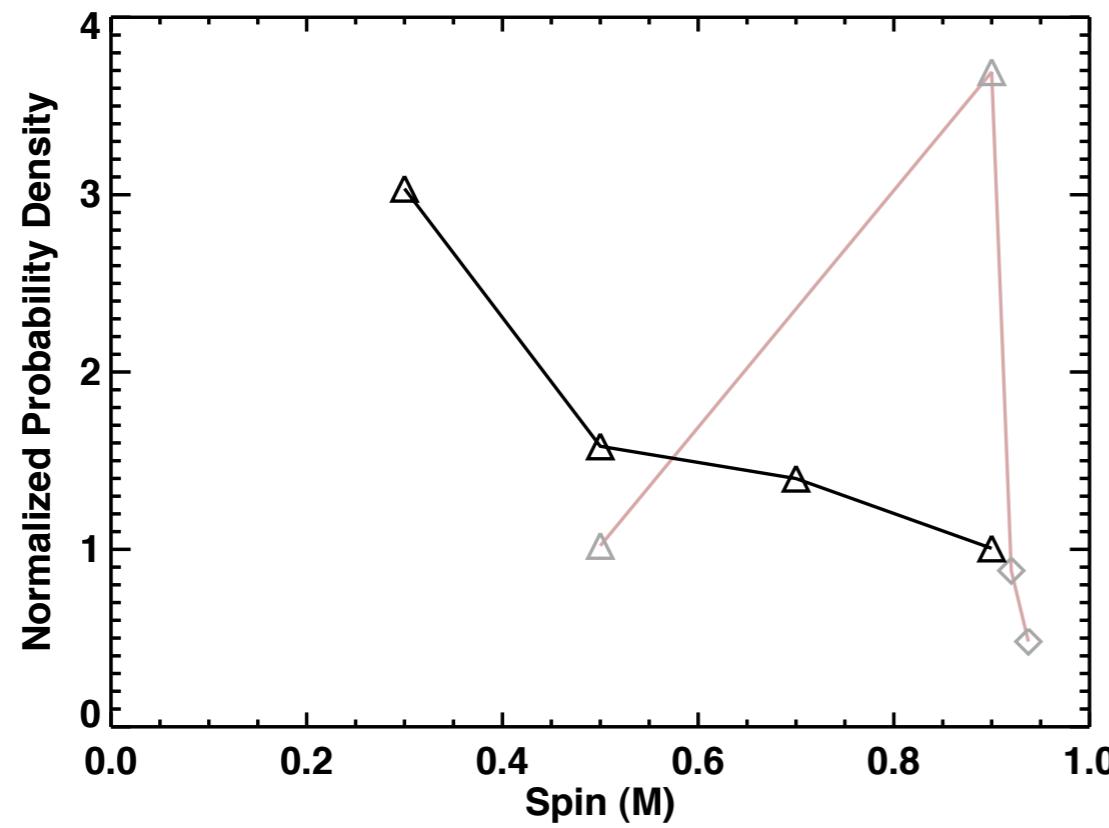
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(Dexter et al. 2010)

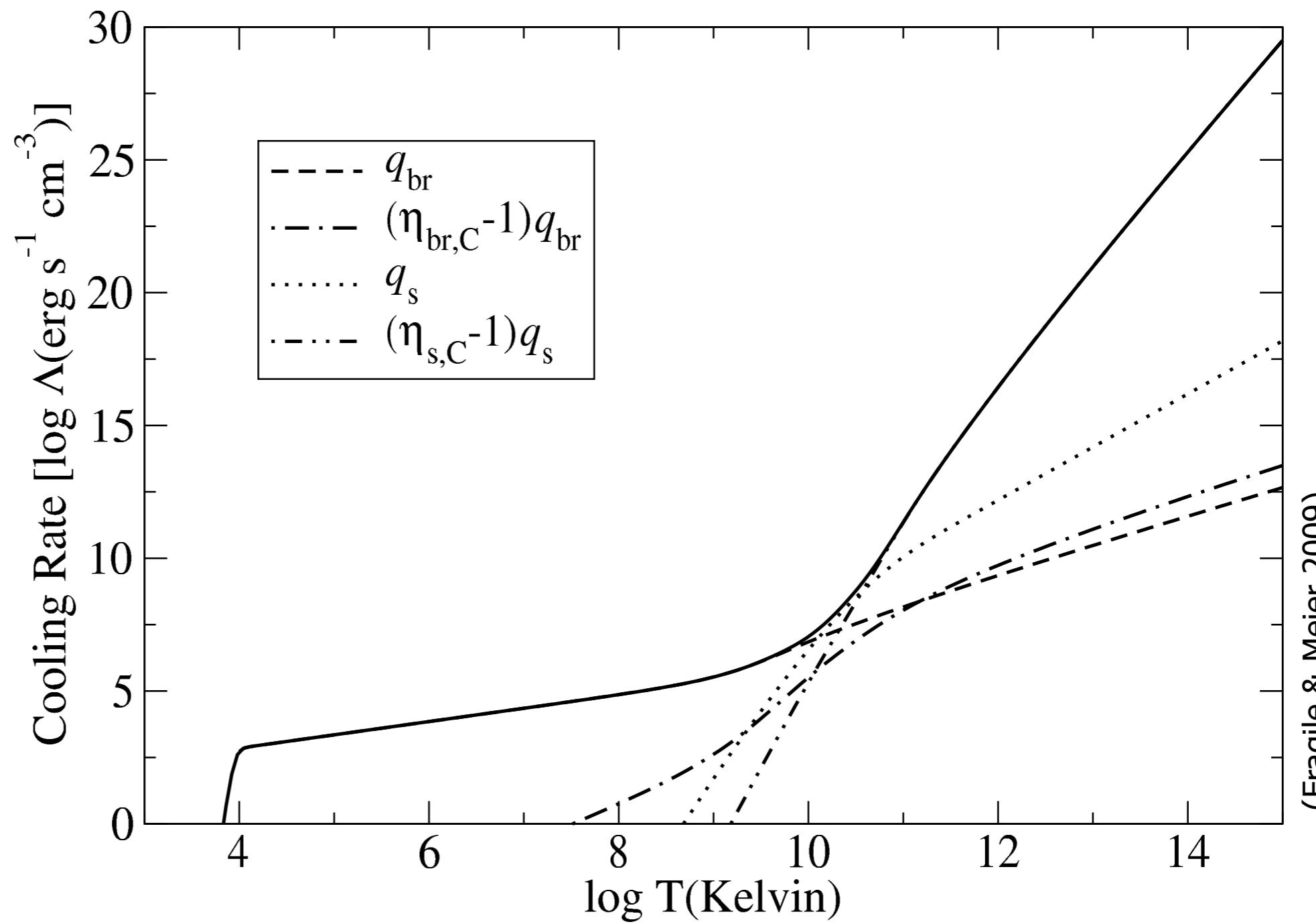


(Dexter & Fragile 2013)

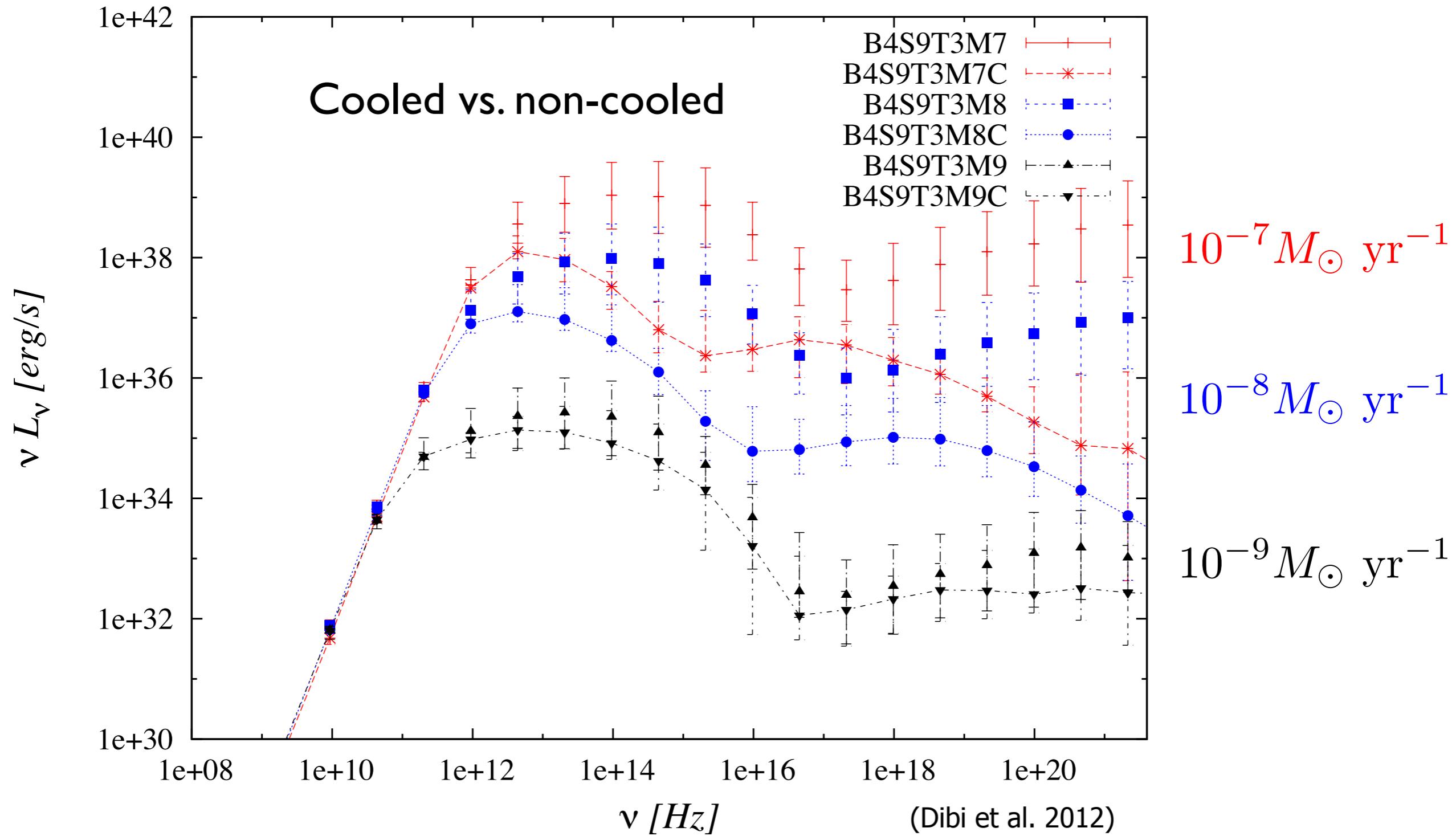
- Cooling simply enters as source term in energy and momentum equations

bremsstrahlung
synchrotron
inverse Compton

$\Lambda(\rho, T, B^2, H)$



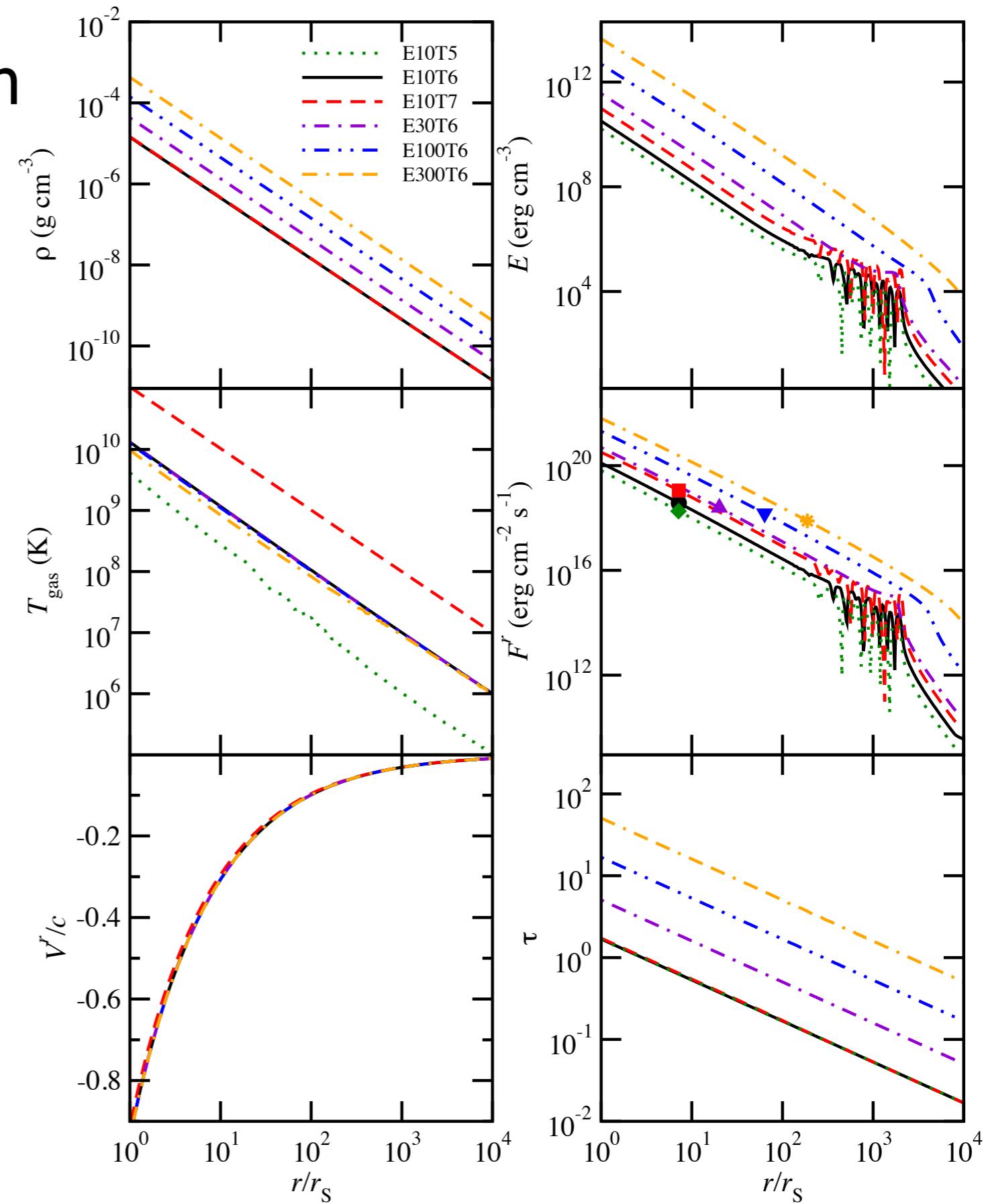
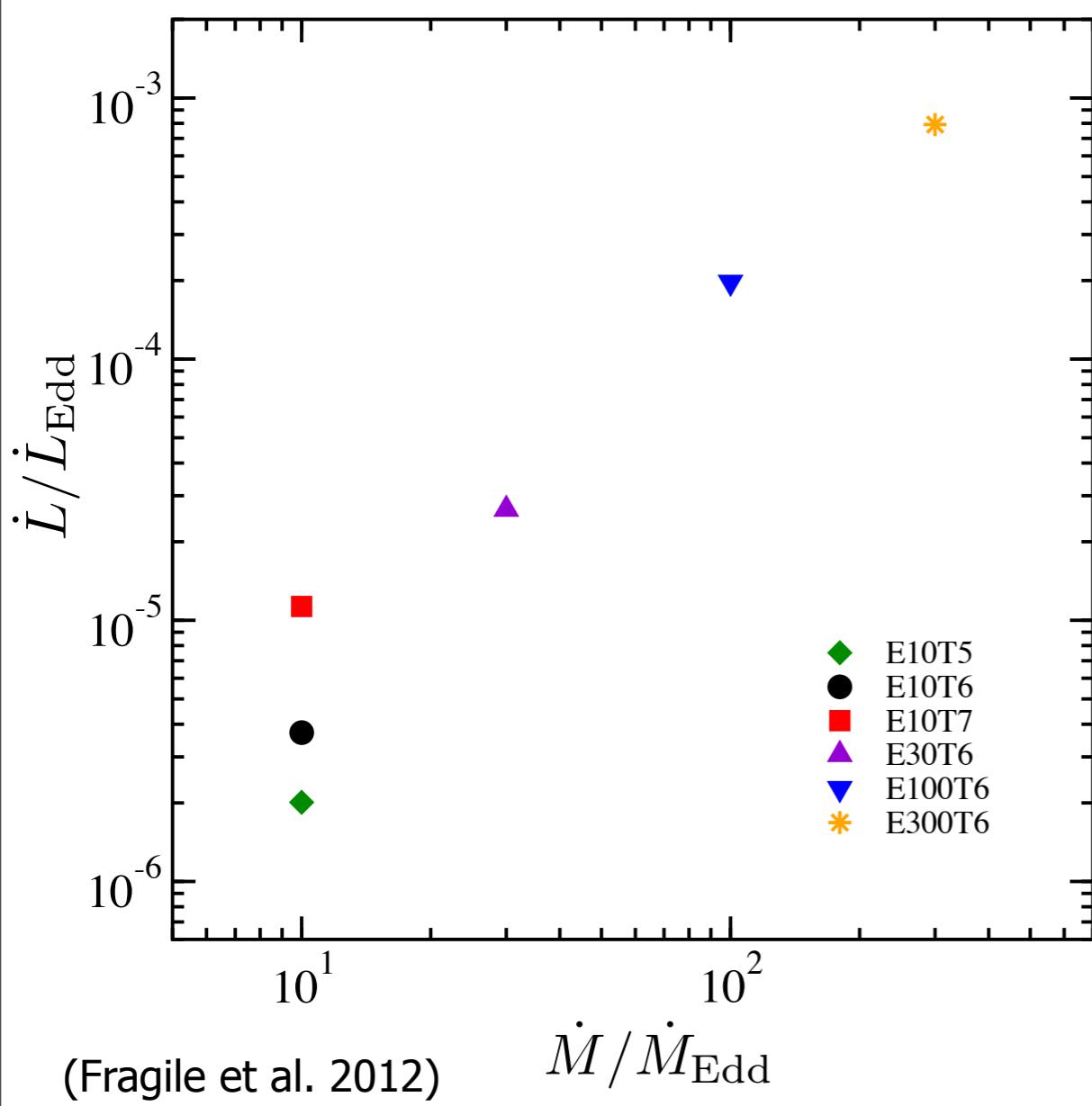
- Is “after-the-fact” radiative treatment adequate for Sgr A*?
 - Dibi et al. (2012); Drappeau et al. (2013)



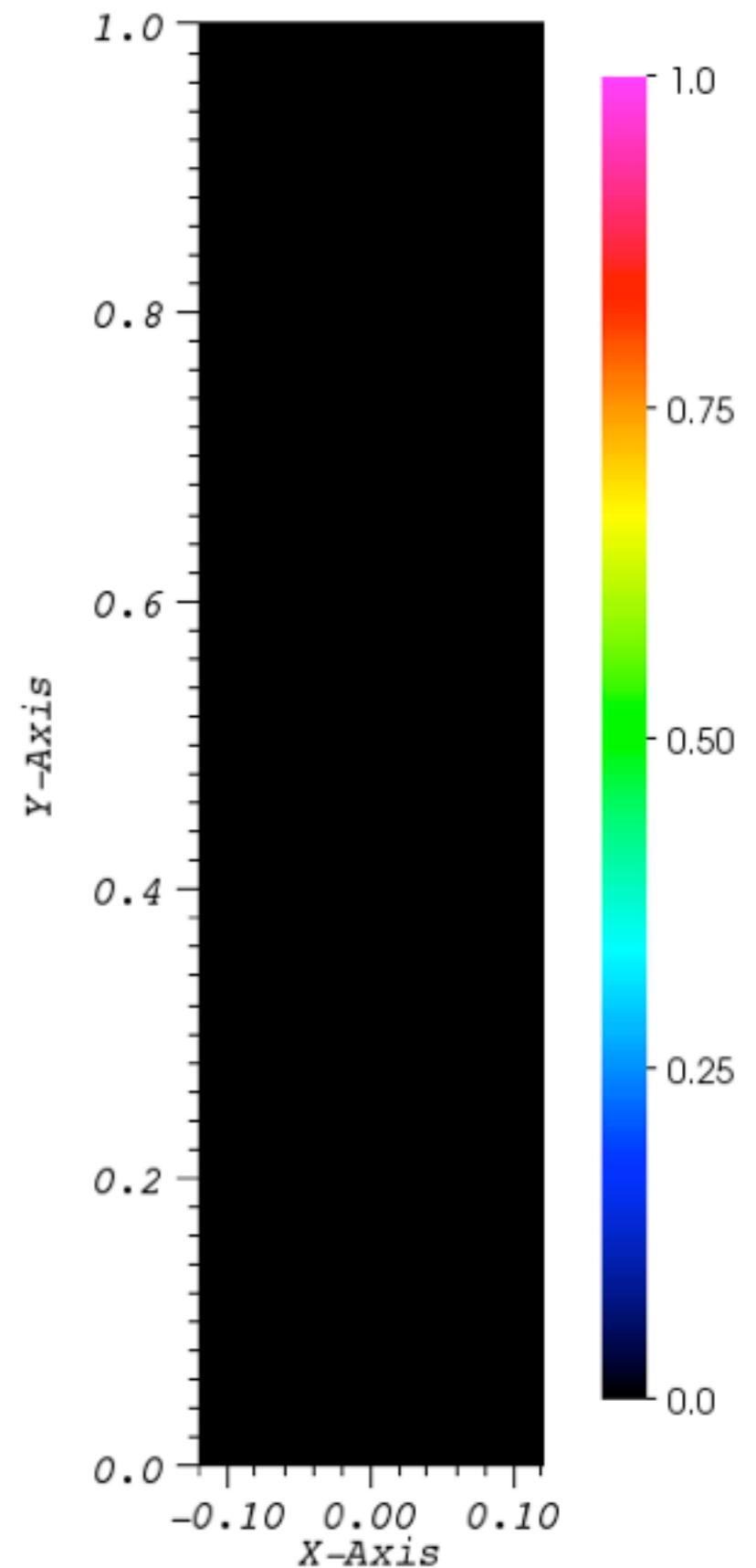
GR Radiation MHD

Step 3: Optically-thick GR rad-MHD

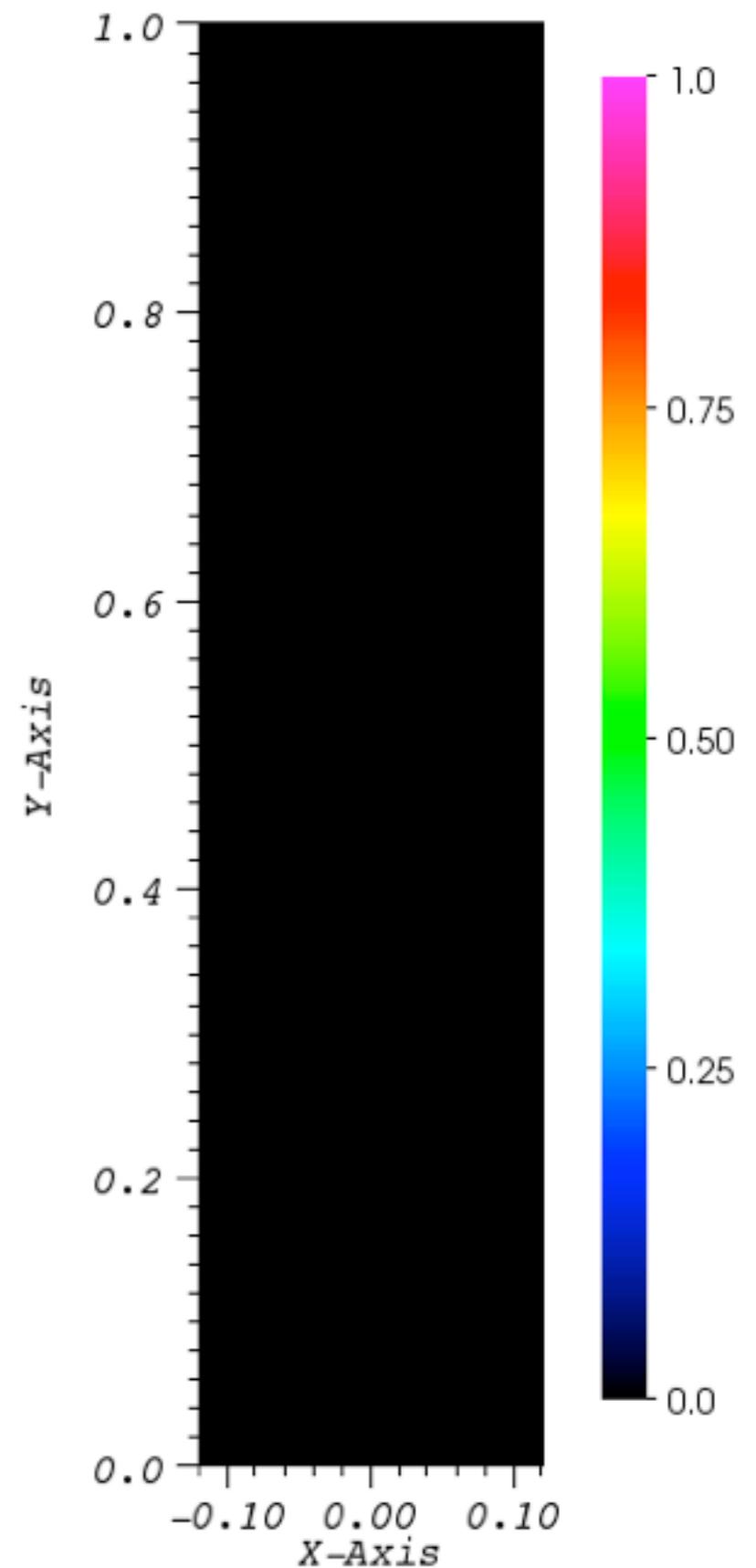
- Bondi inflow with radiation
 - Spherical accretion above Eddington limit



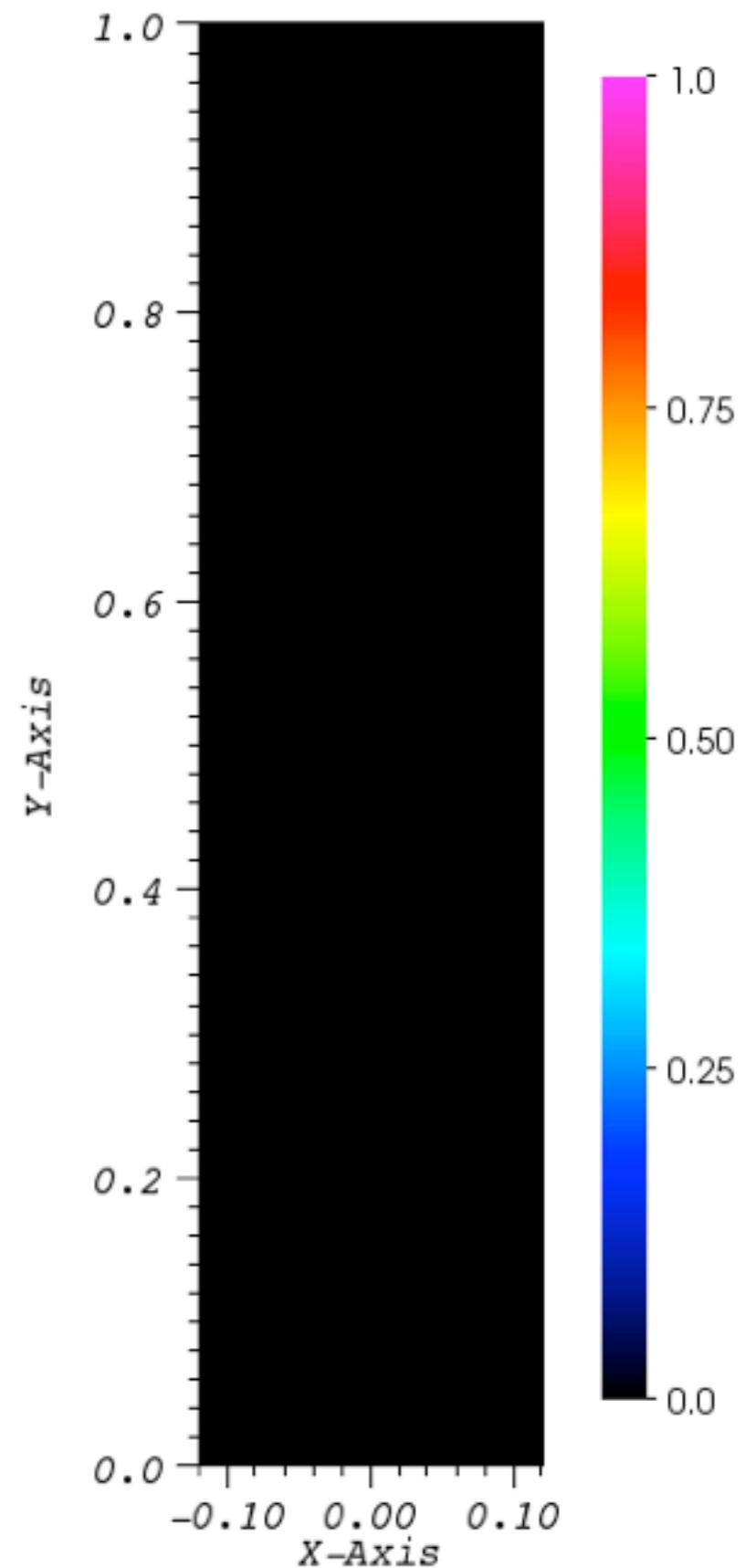
- Drawbacks of method:



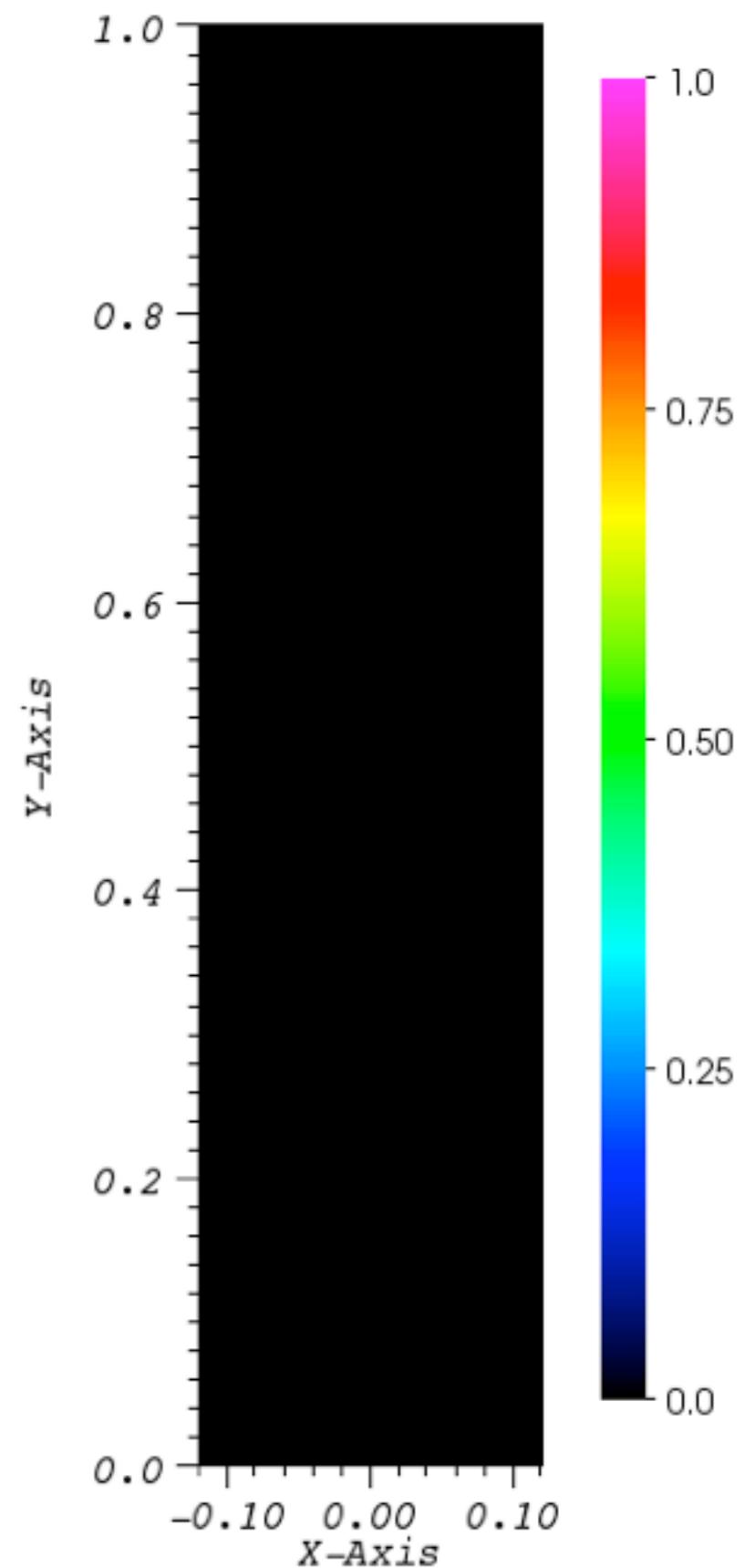
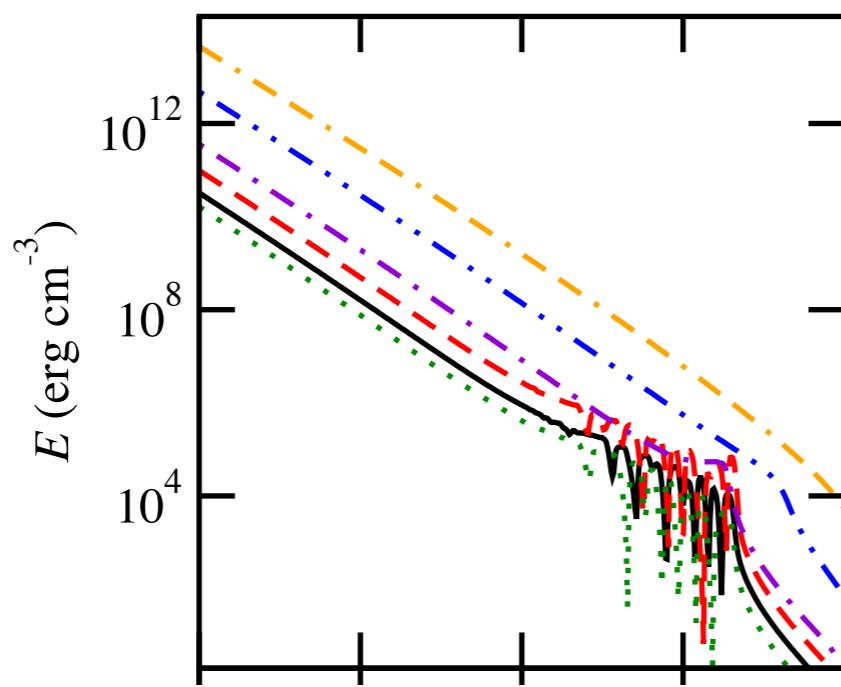
- Drawbacks of method:
 - Eddington closure does not preserve shadows



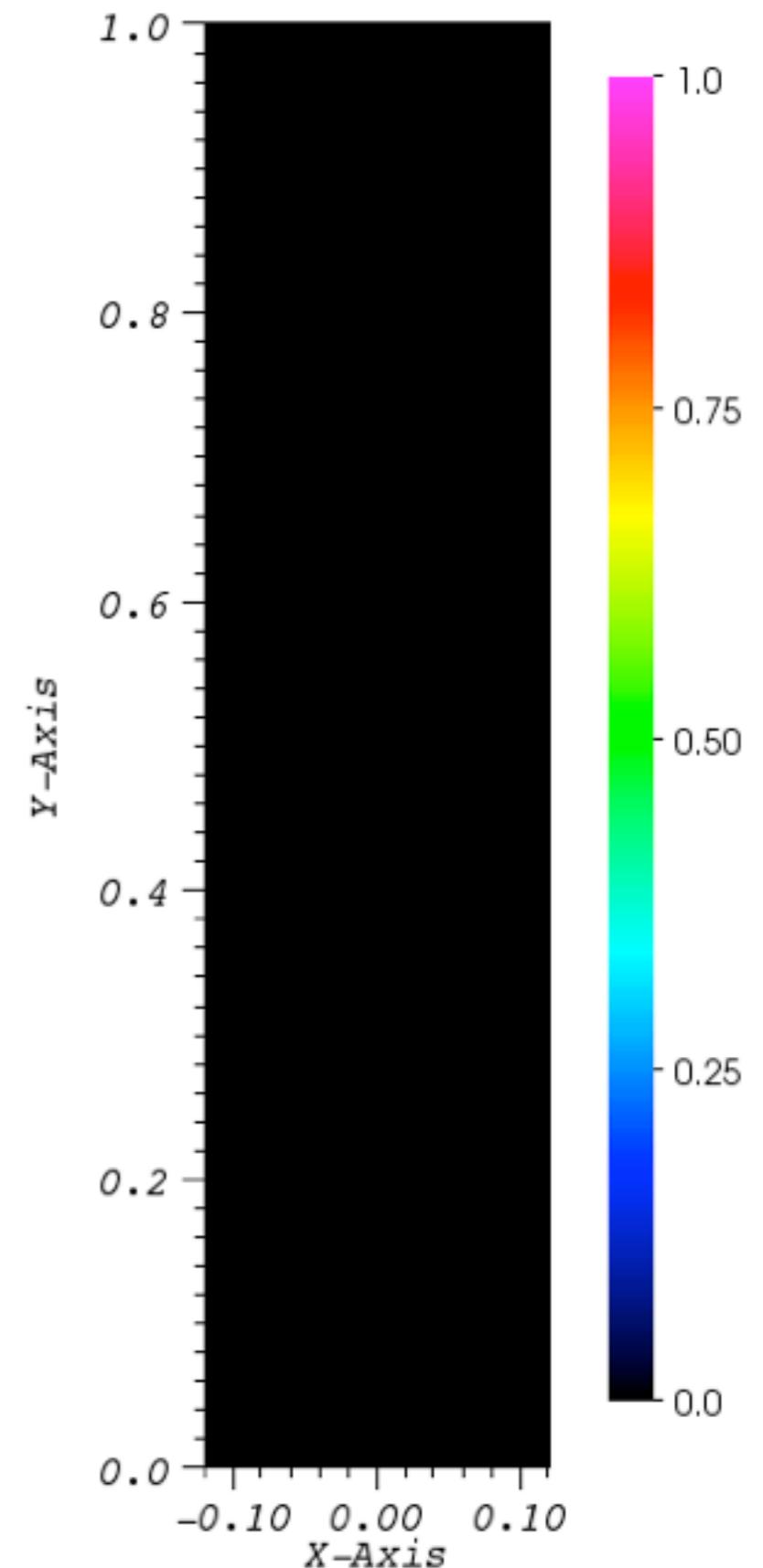
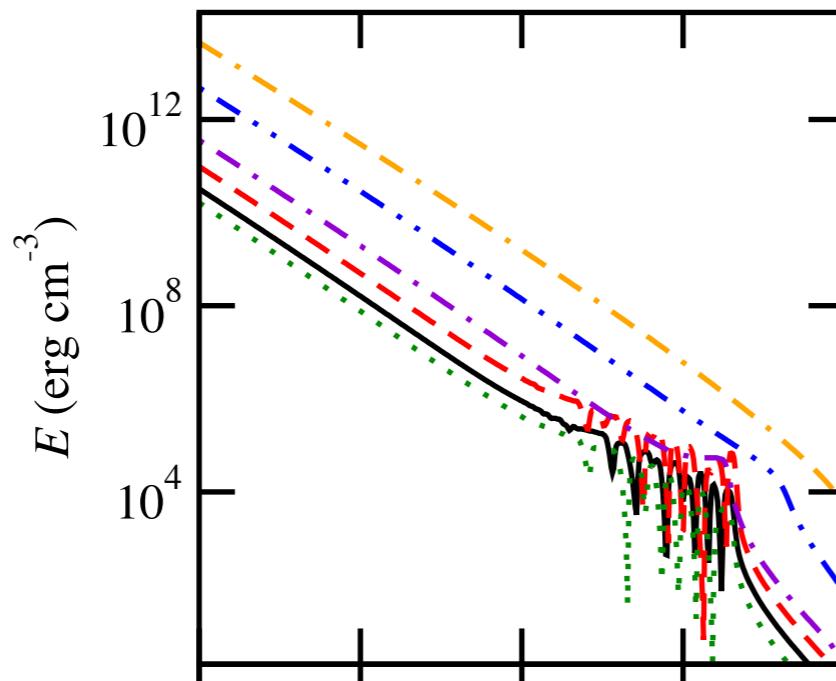
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- Drawbacks of method:
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 - Method becomes unstable in optically thin regions



- Drawbacks of method:
 - Eddington closure does not preserve shadows
 - Method becomes unstable in optically thin regions
 - Radiation source term can be extremely stiff
 - prohibitively small time stepping required by explicit integration



- More general closure (Levermore 1984)

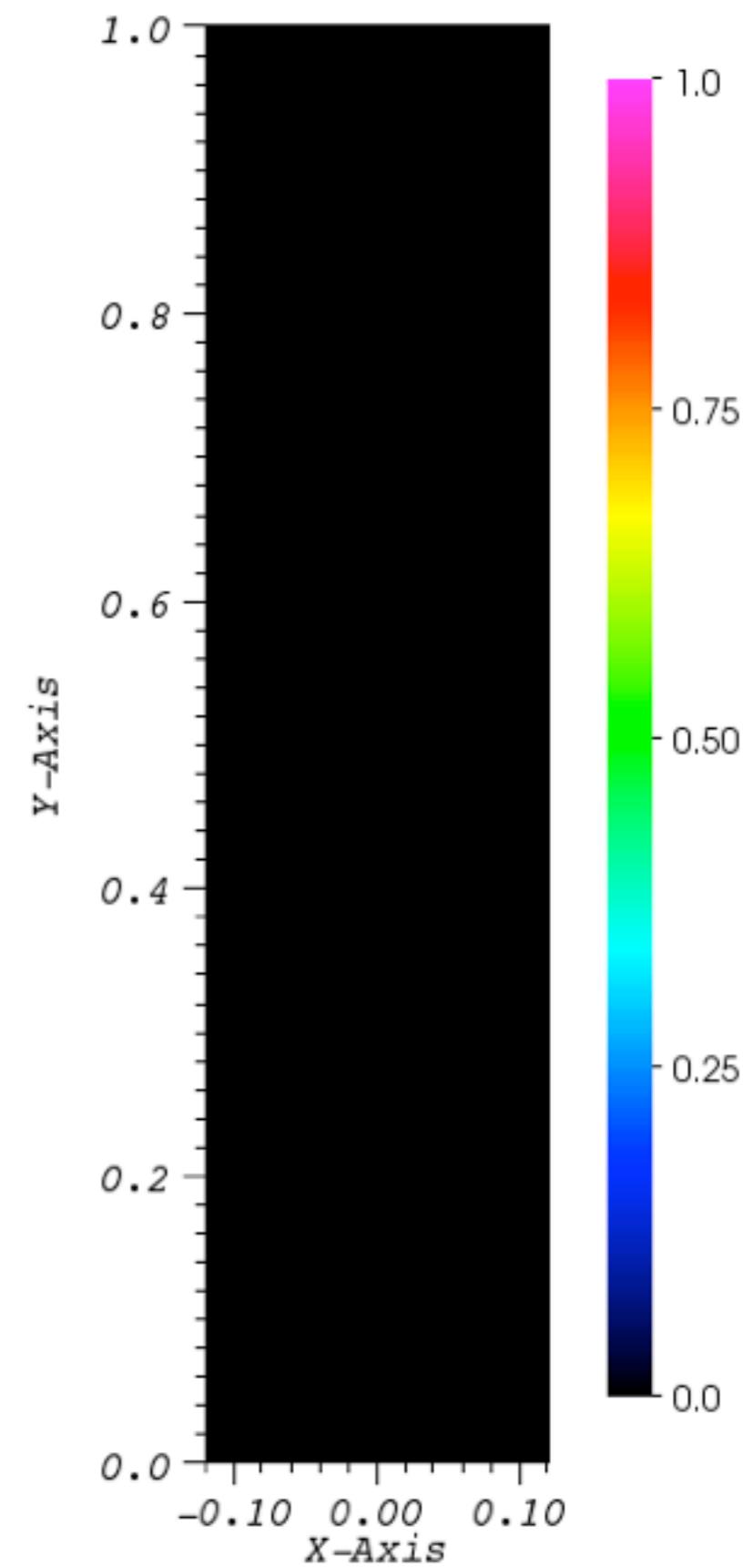
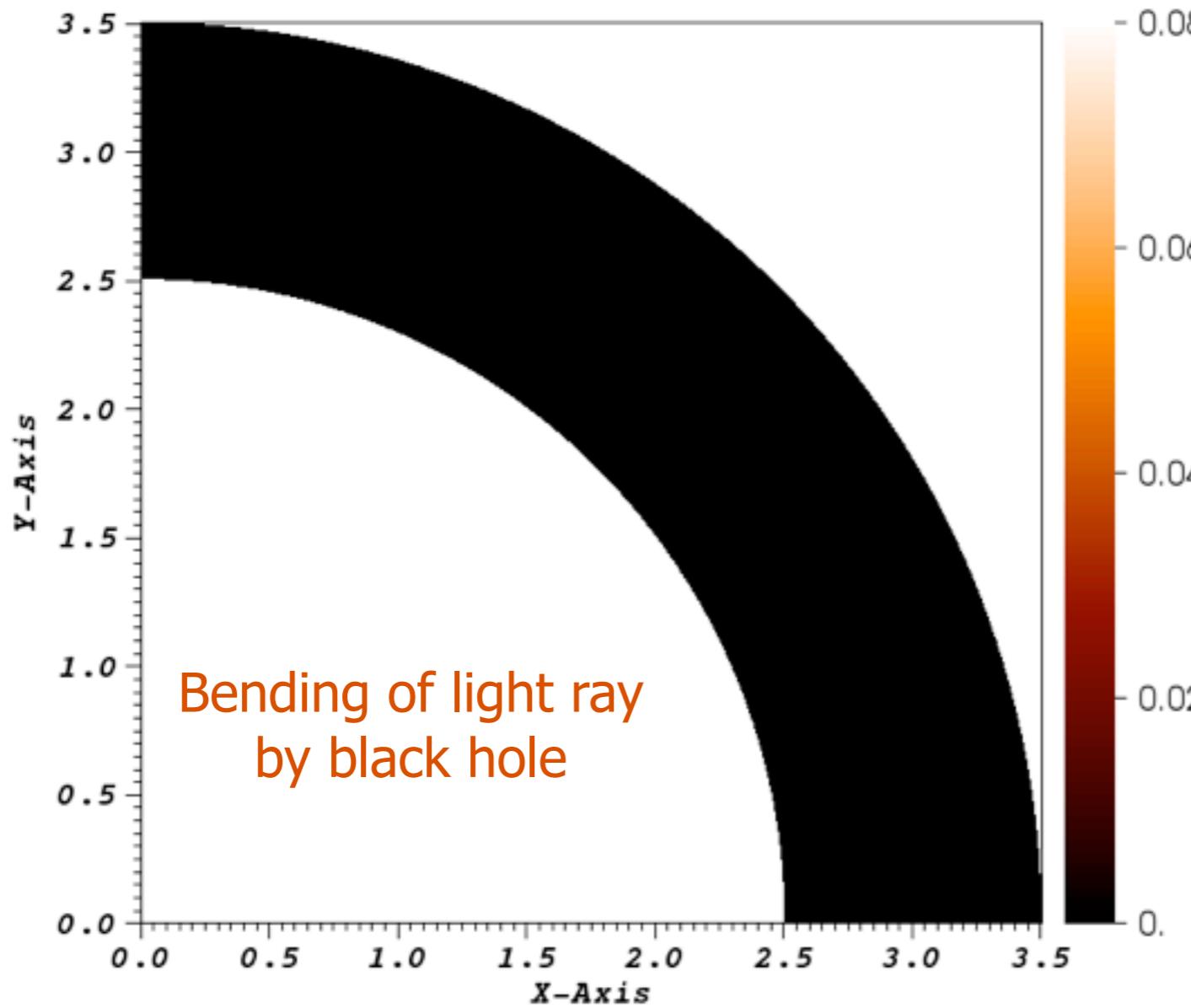
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- Semi-implicit time integration

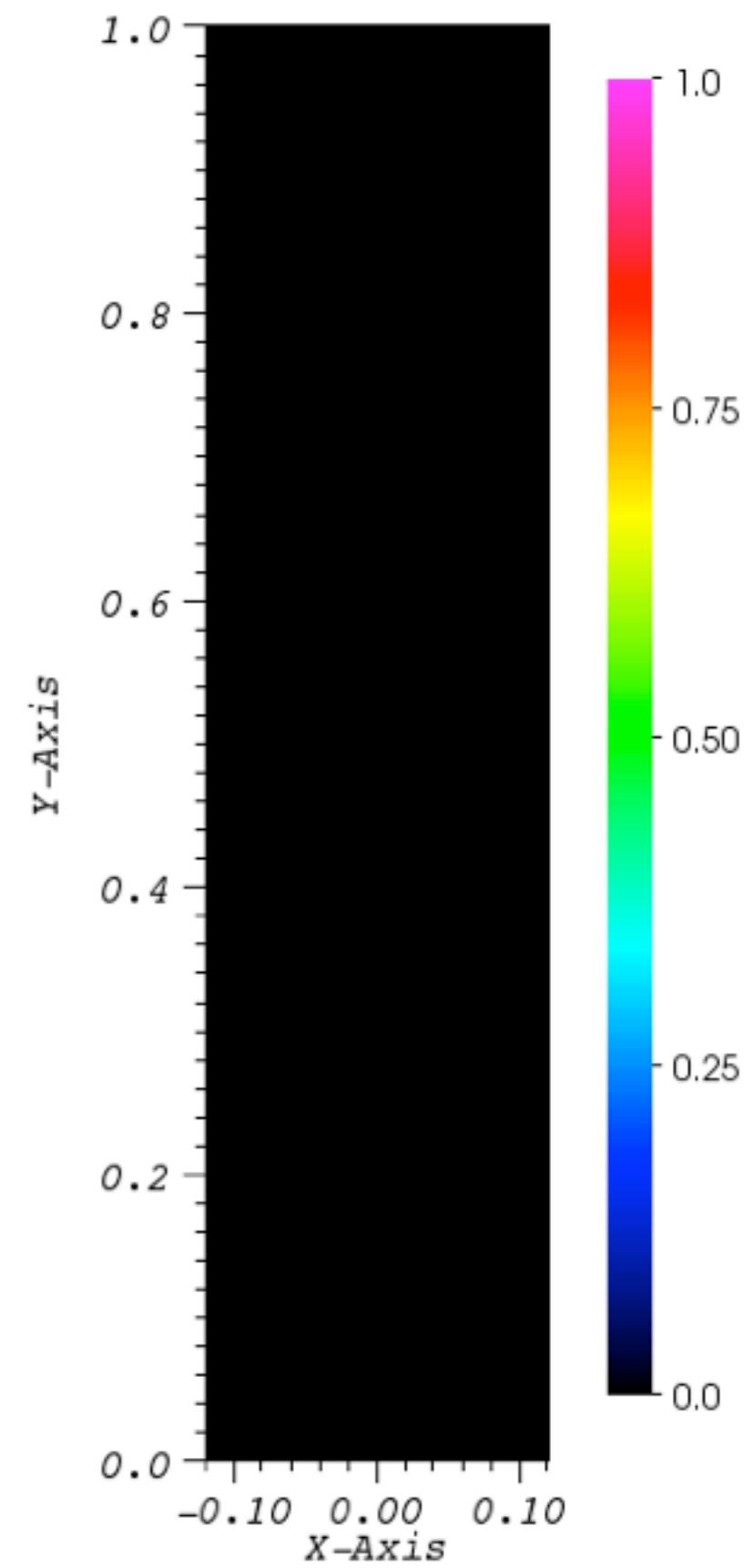
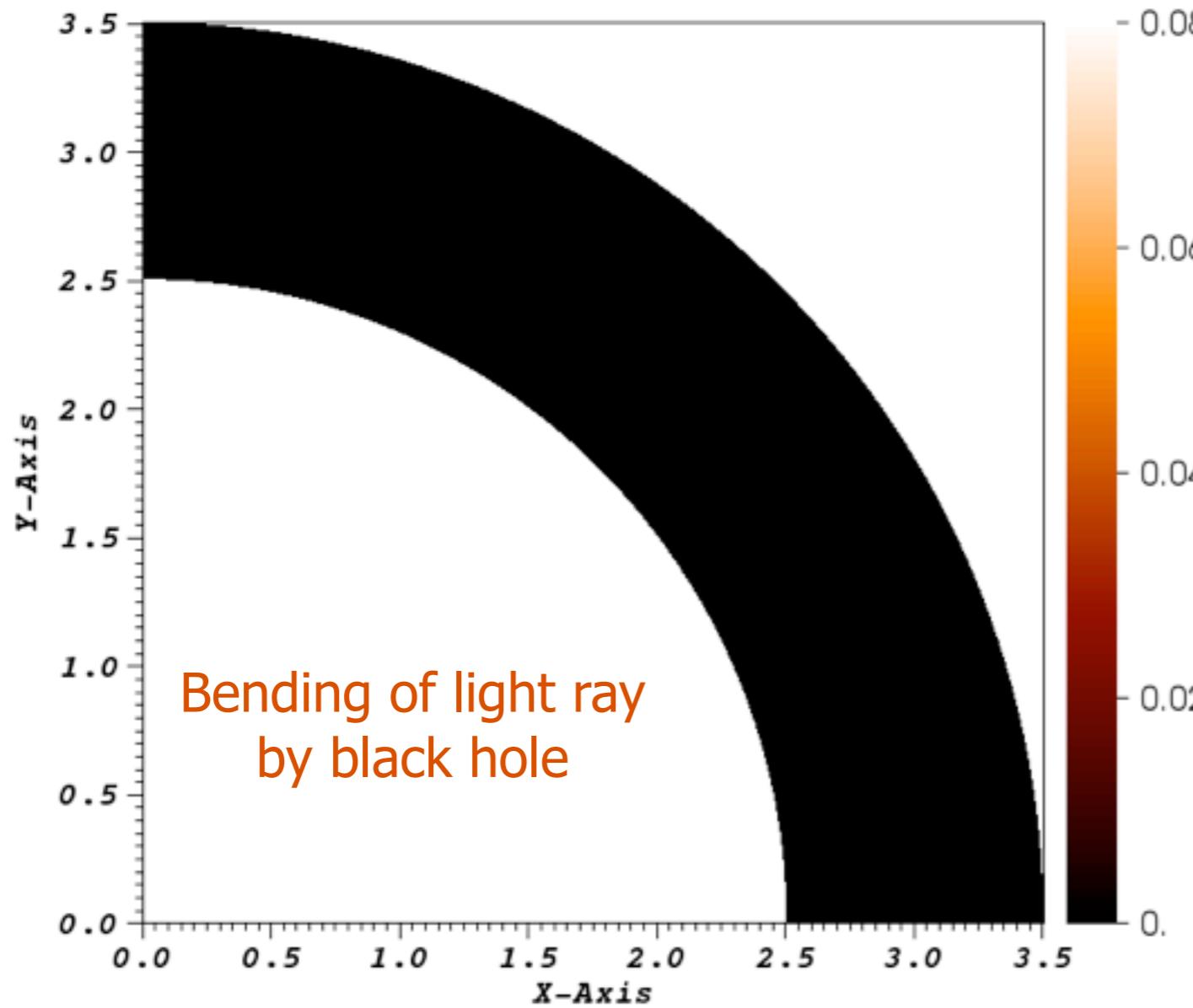
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Bending of light ray
by black hole

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- Thank you
 - KITP
 - program organizers
 - sponsors



South Carolina

